

GALEX UV COLOR RELATIONS FOR NEARBY EARLY-TYPE GALAXIES

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ABSTRACT

We use *GALEX*/optical photometry to construct color-color relationships for early-type galaxies sorted by morphological type. We have matched objects in the *GALEX* GR1 public release and the first IR1.1 internal release, with the RC3 early-type galaxies having a morphological type $-5.5 \leq T < -1.5$, with mean error on $T < 1.5$ and mean error on $(B - V)_T < 0.05$. After visual inspection of each match, we are left with 130 galaxies with reliable *GALEX* pipeline photometry in the far-UV and near-UV bands. This sample is divided into ellipticals ($-5.5 \leq T < -3.5$) and lenticulars ($-3.5 \leq T < -1.5$). After correction for Galactic extinction, the color-color diagrams FUV – NUV versus $(B - V)_{Tc}$ are plotted for the two subsamples. We find a tight anticorrelation between the FUV – NUV and $(B - V)_{Tc}$ colors for ellipticals, with the UV color getting bluer when the $(B - V)_{Tc}$ gets redder. This relationship very likely is an extension of the color-metallicity relationship in the *GALEX* NUV band. We suspect that the main source of the correlation is metal line blanketing in the NUV band. The FUV – NUV versus $B - V$ correlation has larger scatter for lenticular galaxies; we speculate that this reflects the presence of low-level star formation. If the latter objects (i.e., those that are blue both in FUV – NUV and in $B - V$) are interpreted as harboring recent star formation activity, this would be the case for a few percent ($\sim 4\%$) of ellipticals and $\sim 15\%$ of lenticulars; this would mean about 10% of early-type galaxies have residual star formation in our full sample of 130 early-type galaxies. We also plot FUV – NUV versus the Mg_2 index and central velocity dispersion. We find a tight anticorrelation between FUV – NUV and the Mg_2 index; we suspect that this reflects blanketing in the NUV band being correlated with overall metallicity. We find a marginal anticorrelation of FUV – V_T with Mg_2 for elliptical galaxies.

Subject headings: galaxies: elliptical and lenticular, cD — galaxies: photometry — galaxies: stellar content — ultraviolet: galaxies

1. INTRODUCTION

The far-UV radiation from early-type galaxies, first detected by the *Orbiting Astronomical Observatory-2* in 1969 (Code & Welch 1979), is now thought to be produced mainly by low-mass, helium-burning stars in extreme horizontal-branch (EHB) and subsequent phases of evolution. O’Connell (1999) has extensively reviewed the data and the theoretical ideas that contributed to this interpretation.

This minority population of evolved stars has attracted attention beyond the domain of UV observations for several reasons. First, their characterization sheds light on poorly known phases of stellar evolution and associated astrophysical processes (giant-branch mass loss, helium enrichment) that are otherwise difficult to investigate. Second, their evolution has to be understood for more precise modeling of old stellar populations (Charlot et al. 1996) and for using the look-back time of their appearance as constraints on the age of elliptical galaxies (e.g., Greggio & Renzini 1990; Bressan et al. 1996; Yi et al. 1999; Lee et al. 2005). Last, because they are hot, these evolved stars may hide young stars that may trace departure from purely passive evolution, a long-sought feature in the discussion of monolithic versus hierarchical scenarios of galaxy formation.

The most recent advances on the observational front have come from UV images with the *HST* that, at the distance of M32 and the bulge of M31, resolved some of the UV emission into bright stars (Brown et al. 1998, 2000). Recently, a deep UV color-magnitude diagram with the *HST* STIS has confirmed a well-populated EHB in M32 but found fewer post-AGB stars than would have been expected in that population (Brown 2004), revealing a possible discrepancy with current evolution theory. Albeit extremely powerful, such direct observations of individual stars are limited to very few objects and will stay so for the foreseeable future.

A large sample of early-type galaxies is reachable only through integrated light measurements. Following the early observations reviewed by O’Connell (1999), such a possibility is now offered

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by the extensive survey made with the *GALEX* mission (Martin et al. 2005). This survey provides two UV broadband measurements, the far-UV at $\lambda_e = 1528 \text{ \AA}$ (FUV) and the near-UV at $\lambda_e = 2271 \text{ \AA}$ (NUV). The importance of having a NUV band like the NUV band of *GALEX* has been addressed by Fanelli et al. (1988) and Dorman et al. (2003) and its potential for revealing a young stellar component illustrated by Ferreras & Silk (2000) with *HST* observations of a cluster at $z = 0.41$. In a galaxy with no current star formation, the NUV light near 2000 \AA is dominated largely by turnoff and subgiant stars, with some radiation from blue horizontal-branch (HB) stars if present. The FUV radiation at 1500 \AA and shortward is dominated by hot evolved stars, but the nature of these stars remains unsettled. The UV light is of great interest because the UV wavelengths are most sensitive to trace “residual” star formation in elliptical galaxies. In addition to being sensitive to residual star formation at the present epoch, the NUV band is particularly sensitive to the normal passive evolution of a stellar population with look-back time. Color-color diagrams are therefore powerful diagnostics of evolution, residual star formation, and UV-bright populations of evolved stars.

UV observations of elliptical galaxies with *GALEX* have already been reported. Lee et al. (2005) have addressed the issue of observing the UV emission in clusters at different redshifts. Rich et al. (2005) and Yi et al. (2005) have studied samples of early-type galaxies where the *GALEX* photometry has been matched with galaxies in the SDSS main galaxy sample. The first paper concentrates on the FUV–optical color in quiescent early-type galaxies for studying the population responsible for the UV rising flux; the second paper concentrates on the NUV color-magnitude relation, the scatter of which is interpreted in terms of residual star formation (see also Kaviraj et al. 2007; Schawinski et al. 2007). Boselli et al. (2005) have studied a sample of early-type galaxies with a large range of luminosity in the Virgo Cluster. Last, early-type galaxies are included in the *GALEX* Nearby Galaxy Survey (NGS) of Gil de Paz et al. (2007).

Following the completion of the *GALEX* survey over a significant fraction of the sky, this paper can now present the UV observations of nearby early-type galaxies (elliptical/lenticular), with a morphological type selection based on the Third Reference Catalogue of Bright Galaxies (RC3; de Vaucouleurs et al. 1991) and an emphasis on UV-optical color-color diagrams.

2. DATA

2.1. *GALEX* Observations

The *Galaxy Evolution Explorer* (*GALEX*) is a NASA Small Explorer mission launched by a Pegasus-XL vehicle on 2003 April 28 into a 690 km circular orbit. The instrument is designed to image a 1.2° diameter circular field of view with its resolution of $\approx 4.5''$ (FUV), $6.0''$ (NUV) set by detector electronics and to perform imaging or spectroscopy in two UV bands simultaneously, FUV ($1344\text{--}1786 \text{ \AA}$) and NUV ($1771\text{--}2831 \text{ \AA}$). The imaging surveys include an all-sky imaging survey (AIS; $m_{AB} \simeq 20.5$), a medium imaging survey of 1000 deg^2 (MIS; $m_{AB} \simeq 23.0$), a deep imaging survey of 100 deg^2 (DIS; $m_{AB} \simeq 25.0$), and a NGS. The *GALEX* instrument and mission are described by Martin et al. (2005) and Morrissey et al. (2005). All observations are performed in a pointed mode with an arcminute spiral dither. From the time-tagged photon lists, the data analysis pipeline generates integrated images. The data pipeline includes a photometric module which uses the SExtractor program (Bertin & Arnouts 1996) for detection and photometry of the sources in the imaging data, with some modifications in the determination of sky background.

The *GALEX* pipeline makes a few different measurements of the UV flux of a source. In the following we use the UV flux measured through seven circular apertures with fixed diameters ($3.0''$, $4.5''$, $7.5''$, $12.0''$, $18.0''$, $25.5''$, and $34.5''$) and two elliptical apertures whose major-axis D_{NUV} and D_{FUV} are scaled to 2.5 times the Kron diameter (Kron 1980) determined in the NUV and FUV images, respectively.

GALEX uses the AB magnitude system of Oke & Gunn (1983), with the FUV and NUV magnitude defined by

$$m_{UV} = m_0 - 2.5 \log f_{UV},$$

where f_{UV} is the count rate (flat-field-corrected) and m_0 is the zero point ($m_0 = 20.082$ and 18.817 for NUV and FUV, respectively) corresponding to the AB magnitude of 1 count s^{-1} flat-field-corrected detection. The uncertainty on the *GALEX* zero points is estimated to be ± 0.15 mag for both the FUV and NUV channels.

2.2. Morphological Selection in the RC3 Catalog

The RC3 catalog is reported to be reasonably complete for galaxies having apparent diameters larger than $1'$ at the D_{25} isophotal level¹³ and total B -band magnitudes B_T brighter than about 15.5 , with a redshift not in excess of $15,000 \text{ km s}^{-1}$. Additional objects meeting only the diameter or the magnitude condition and objects of interest smaller than $1'$, fainter than 15.5 , or with redshifts $> 15,000 \text{ km s}^{-1}$ are also included in the RC3. The total number of RC3 objects is 23,022; of them, 1103 have $(B - V)_T$ measurements¹⁴ and are classified as ellipticals or lenticulars/S0s, which correspond to values of the numerical index $-5.5 \leq T < -0.5$.

In order to obtain a reliable subsample of early-type galaxies, we select only the sources with morphological types in the range $-5.5 \leq T < -1.5$, with mean error on $T < 1.5$ and mean error on $(B - V)_T < 0.05$; this leads to an RC3 subsample of 875 galaxies.

2.3. Resulting Sample

In the RC3 subsample of 875 early-type galaxies, 138 (i.e., $\sim 16\%$) are in the fields observed in the FUV and NUV bands by *GALEX* in the AIS, MIS, and NGS included in the GR1 public release and the first IR1.1 internal release.

Because the coordinates in the RC3 are not very accurate, we have searched for UV counterparts within a radius of $1.5'$ and examined individually each field in the FUV, NUV, and POSS I E images; this is possible because the objects are relatively bright and small in number. After visual inspection we find 130 galaxies with a reliable *GALEX* pipeline photometry; 6 have doubtful photometry because of serious blend with close sources (PGC 42620, 69260, 26377, 2569, 13457, and 14375), and 2 are not detected in the FUV because of short exposures (PGC 65580 and 68241). Since 130 of the 138 galaxies have reliable UV photometry, we conclude that our sample is not UV magnitude-limited and could indeed be representative of the RC3 magnitude- and diameter-limited sample. The 130 early-type galaxies of our resulting sample are listed in Table 1 with some relevant information.

¹³ D_{25} is the apparent major isophotal diameter measured at or reduced to surface brightness level $\mu_B = 25.0$.

¹⁴ $(B - V)_T$ is the total (asymptotic) color index in the Johnson $B - V$ system.

TABLE 1
THE EARLY-TYPE GALAXIES SAMPLE

PGC (1)	Alt. Name (2)	T (3)	$E(B-V)$ (4)	$(B-V)_{\text{TC}}$ (5)	$\log D_{25}$ (6)	$\log D_{\text{NUV}}$ (7)	NUV (8)	FUV (9)	$FUV - V_T$ (10)	Mg ₂ (11)	$\log \sigma_V$ (12)
PGC 34780	NGC 3641	-5.3 ± 0.4	0.042	0.74 ± 0.02	1.03	1.41	17.59 ± 0.009	19.13 ± 0.042	6.10 ± 0.21	0.28 ± 0.006	2.21 ± 0.010
PGC 02149	NGC 163	-5.0 ± 0.4	0.034	0.79 ± 0.02	1.19	1.22	18.15 ± 0.013	19.28 ± 0.037	6.73 ± 0.15	...	2.31 ± 0.026
PGC 03513	UGC 610	-5.0 ± 0.8	0.071	0.86 ± 0.04	1.26	1.08	19.18 ± 0.151	20.05 ± 0.218	7.36 ± 0.30
PGC 04376	NGC 430	-5.0 ± 0.6	0.030	0.83 ± 0.03	1.12	0.97	18.27 ± 0.067	19.54 ± 0.157	7.16 ± 0.43	0.28 ± 0.009	2.46 ± 0.026
PGC 05663	NGC 584	-5.0 ± 0.3	0.042	0.80 ± 0.01	1.62	1.58	15.81 ± 0.004	17.63 ± 0.019	7.33 ± 0.13	0.26 ± 0.003	2.32 ± 0.009
PGC 07584	NGC 777	-5.0 ± 0.3	0.047	0.87 ± 0.02	1.39	1.29	17.19 ± 0.008	17.70 ± 0.020	6.44 ± 0.14	0.33 ± 0.008	2.51 ± 0.014
PGC 08557	NGC 855	-5.0 ± 0.7	0.072	0.52 ± 0.02	1.42	1.38	15.31 ± 0.004	15.93 ± 0.010	5.31 ± 0.13	0.03 ± 0.008	1.85 ± 0.058
PGC 09890	NGC 990	-5.0 ± 0.8	0.112	0.80 ± 0.02	1.26	1.10	17.70 ± 0.073	18.90 ± 0.139	6.84 ± 0.21	0.24 ± 0.005	2.22 ± 0.030
PGC 10175	NGC 1052	-5.0 ± 0.3	0.027	0.79 ± 0.01	1.48	1.58	15.83 ± 0.003	16.85 ± 0.010	6.51 ± 0.13	0.30 ± 0.005	2.32 ± 0.008
PGC 13299	NGC 1379	-5.0 ± 0.4	0.012	0.76 ± 0.01	1.38	1.56	16.27 ± 0.004	18.05 ± 0.017	7.22 ± 0.10	0.25 ± 0.003	2.08 ± 0.010
PGC 13418	NGC 1399	-5.0 ± 0.3	0.013	0.83 ± 0.01	1.84	1.51	14.72 ± 0.002	15.20 ± 0.004	5.70 ± 0.10	0.33 ± 0.003	2.53 ± 0.007
PGC 13419	NGC 1395	-5.0 ± 0.3	0.023	0.82 ± 0.01	1.77	1.47	15.62 ± 0.017	16.44 ± 0.037	6.96 ± 0.07	0.31 ± 0.003	2.39 ± 0.009
PGC 13433	NGC 1404	-5.0 ± 0.3	0.011	0.84 ± 0.01	1.52	1.61	15.36 ± 0.003	16.41 ± 0.009	6.49 ± 0.13	0.31 ± 0.003	2.37 ± 0.008
PGC 13505	NGC 1407	-5.0 ± 0.3	0.069	0.84 ± 0.01	1.66	1.74	14.93 ± 0.003	15.64 ± 0.010	6.22 ± 0.20	0.32 ± 0.004	2.44 ± 0.009
PGC 13638	NGC 1426	-5.0 ± 0.3	0.017	0.76 ± 0.01	1.42	1.32	16.84 ± 0.031	18.30 ± 0.083	7.01 ± 0.10	0.24 ± 0.006	2.19 ± 0.009
PGC 13738	NGC 1439	-5.0 ± 0.3	0.030	0.73 ± 0.01	1.39	1.33	16.74 ± 0.033	18.44 ± 0.094	7.18 ± 0.11	0.27 ± 0.009	...
PGC 13814	NGC 1453	-5.0 ± 0.4	0.105	0.83 ± 0.01	1.38	1.15	17.03 ± 0.062	18.16 ± 0.117	7.00 ± 0.18	0.30 ± 0.004	2.48 ± 0.032
PGC 14520	NGC 1521	-5.0 ± 0.5	0.040	0.81 ± 0.01	1.44	1.31	17.15 ± 0.033	18.66 ± 0.110	7.41 ± 0.13	0.26 ± 0.005	2.38 ± 0.019
PGC 14757	NGC 1549	-5.0 ± 0.6	0.013	0.80 ± 0.01	1.69	1.90	14.83 ± 0.002	16.55 ± 0.013	6.84 ± 0.08	0.28 ± 0.016	2.31 ± 0.008
PGC 22322	NGC 2475	-5.0 ± 0.9	0.039	0.79 ± 0.03	0.90	0.95	18.69 ± 0.061	19.63 ± 0.154
PGC 24909	NGC 2675	-5.0 ± 0.8	0.020	0.90 ± 0.03	1.18	1.24	18.84 ± 0.016	20.39 ± 0.057	7.24 ± 0.21
PGC 25144	NGC 2693	-5.0 ± 0.5	0.021	0.82 ± 0.01	1.42	1.35	17.15 ± 0.009	18.04 ± 0.023	6.27 ± 0.15	0.32 ± 0.004	2.41 ± 0.047
PGC 26514	NGC 2810	-5.0 ± 0.8	0.052	0.84 ± 0.02	1.23	1.03	18.14 ± 0.061	18.80 ± 0.116	6.77 ± 0.19	0.30 ± 0.006	2.54 ± 0.013
PGC 29822	NGC 3158	-5.0 ± 0.5	0.013	0.88 ± 0.01	1.30	1.14	17.79 ± 0.048	18.26 ± 0.080	6.41 ± 0.15	0.32 ± 0.004	2.35 ± 0.014
PGC 29825	NGC 3193	-5.0 ± 0.8	0.013	0.85 ± 0.02	1.32	0.93	19.18 ± 0.099	20.01 ± 0.172	6.46 ± 0.22
PGC 30099	NGC 3159	-5.0 ± 0.3	0.026	0.80 ± 0.01	1.48	1.52	16.14 ± 0.005	17.63 ± 0.018	6.87 ± 0.04	0.26 ± 0.010	2.29 ± 0.013
PGC 34778	NGC 3640	-5.0 ± 0.3	0.044	0.76 ± 0.01	1.60	1.61	15.74 ± 0.004	17.78 ± 0.022	7.52 ± 0.13	0.24 ± 0.005	2.26 ± 0.010
PGC 37061	NGC 3923	-5.0 ± 0.3	0.083	0.80 ± 0.02	1.77	1.88	14.68 ± 0.002	16.10 ± 0.017	6.60 ± 0.40	0.30 ± 0.005	2.40 ± 0.010
PGC 39764	NGC 4278	-5.0 ± 0.3	0.029	0.78 ± 0.01	1.61	1.67	15.08 ± 0.003	16.01 ± 0.009	5.98 ± 0.13	0.27 ± 0.001	2.38 ± 0.010
PGC 39800	NGC 4283	-5.0 ± 0.4	0.025	0.79 ± 0.01	1.18	1.47	17.11 ± 0.008	18.52 ± 0.027	6.57 ± 0.13	0.25 ± 0.004	2.05 ± 0.016
PGC 40455	NGC 4374	-5.0 ± 0.3	0.040	0.82 ± 0.01	1.81	1.88	14.35 ± 0.002	15.60 ± 0.006	6.65 ± 0.05	0.29 ± 0.001	2.45 ± 0.004
PGC 40562	NGC 4387	-5.0 ± 0.6	0.033	0.74 ± 0.01	1.25	1.23	17.44 ± 0.007	19.09 ± 0.027	7.11 ± 0.06	0.23 ± 0.004	2.02 ± 0.011
PGC 40653	NGC 4406	-5.0 ± 0.3	0.030	0.78 ± 0.01	1.95	2.13	14.09 ± 0.002	15.43 ± 0.005	6.67 ± 0.05	0.29 ± 0.003	2.37 ± 0.005
PGC 41095	NGC 4458	-5.0 ± 0.4	0.024	0.72 ± 0.01	1.24	1.39	17.22 ± 0.006	19.38 ± 0.028	7.43 ± 0.05	0.21 ± 0.003	2.01 ± 0.011
PGC 41228	NGC 4473	-5.0 ± 0.3	0.028	0.81 ± 0.01	1.65	1.72	15.43 ± 0.003	16.93 ± 0.009	6.86 ± 0.04	0.30 ± 0.004	2.25 ± 0.007
PGC 41963	NGC 4551	-5.0 ± 0.6	0.039	0.79 ± 0.01	1.26	1.30	17.27 ± 0.006	18.97 ± 0.023	7.11 ± 0.06	0.25 ± 0.003	2.03 ± 0.009
PGC 41968	NGC 4552	-5.0 ± 0.4	0.041	0.82 ± 0.01	1.71	1.78	14.83 ± 0.002	15.73 ± 0.005	6.15 ± 0.05	0.30 ± 0.005	2.40 ± 0.005
PGC 51275	NGC 5576	-5.0 ± 0.3	0.031	0.74 ± 0.01	1.55	1.68	15.60 ± 0.004	17.55 ± 0.018	6.73 ± 0.13	0.24 ± 0.005	2.26 ± 0.016
PGC 51787	NGC 5638	-5.0 ± 0.3	0.033	0.79 ± 0.01	1.43	1.62	16.27 ± 0.005	17.92 ± 0.021	6.86 ± 0.14	0.30 ± 0.004	2.22 ± 0.009
PGC 53643	NGC 5813	-5.0 ± 0.3	0.057	0.81 ± 0.01	1.62	1.61	15.99 ± 0.010	17.10 ± 0.023	6.87 ± 0.13	0.29 ± 0.004	2.38 ± 0.008
PGC 61009	NGC 6482	-5.0 ± 0.6	0.099	0.68 ± 0.01	1.30	1.14	16.29 ± 0.006	18.61 ± 0.046	7.51 ± 0.14	0.32 ± 0.003	2.49 ± 0.015
PGC 67882	NGC 7168	-5.0 ± 0.4	0.023	0.79 ± 0.01	1.31	1.25	17.41 ± 0.043	18.94 ± 0.115	7.12 ± 0.17	...	2.26 ± 0.020
PGC 69256	NGC 7317	-5.0 ± 0.5	0.080	0.77 ± 0.03	1.05	0.88	19.34 ± 0.028	20.05 ± 0.105	6.74 ± 0.16	0.32 ± 0.003	...
PGC 70090	IC 1459	-5.0 ± 0.3	0.016	0.85 ± 0.01	1.72	1.48	15.58 ± 0.017	16.49 ± 0.038	6.59 ± 0.16
PGC 49248	NGC 5329	-4.7 ± 0.6	0.029	0.79 ± 0.03	1.12	1.35	17.91 ± 0.008	19.57 ± 0.039	7.28 ± 0.16
PGC 03342	UGC 579	-4.5 ± 0.6	0.039	0.89 ± 0.01	1.16	1.05	19.08 ± 0.127	19.40 ± 0.150	6.22 ± 0.21	...	2.49 ± 0.042
PGC 05231	IC 1696	-4.3 ± 0.4	0.042	0.86 ± 0.02	0.94	1.04	18.98 ± 0.019	20.46 ± 0.065	7.03 ± 0.16	0.29 ± 0.004	2.23 ± 0.012

TABLE 1—Continued

PGC (1)	Alt. Name (2)	T (3)	$E(B - V)$ (4)	$(B - V)_{TC}$ (5)	$\log D_{25}$ (6)	$\log D_{NUV}$ (7)	NUV (8)	FUV (9)	$FUV - V_T$ (10)	M_{G_2} (11)	$\log \sigma_V$ (12)
PGC 03455	NGC 315	-4.0 ± 0.5	0.065	0.86 ± 0.02	1.51	1.30	16.84 ± 0.021	17.58 ± 0.047	6.66 ± 0.21	0.30 ± 0.004	2.43 ± 0.044
PGC 04224	NGC 410	-4.0 ± 0.6	0.058	0.86 ± 0.01	1.38	1.41	17.08 ± 0.010	17.93 ± 0.029	6.67 ± 0.13	0.34 ± 0.003	2.48 ± 0.010
PGC 04363	NGC 426	-4.0 ± 0.5	0.032	0.80 ± 0.04	1.14	0.81	17.95 ± 0.051	18.67 ± 0.109	5.96 ± 0.42	...	2.44 ± 0.062
PGC 05766	NGC 596	-4.0 ± 0.5	0.058	0.72 ± 0.01	1.51	1.54	15.93 ± 0.005	17.77 ± 0.021	7.06 ± 0.12	0.23 ± 0.003	2.18 ± 0.012
PGC 58265	NGC 6166	-4.0 ± 0.4	0.012	0.87 ± 0.04	1.29	1.51	17.41 ± 0.009	18.12 ± 0.023	6.42 ± 0.12	0.32 ± 0.004	2.50 ± 0.008
PGC 65379	NGC 6964	-4.0 ± 0.5	0.094	0.80 ± 0.01	1.23	1.38	17.52 ± 0.009	18.64 ± 0.054	5.99 ± 0.21	...	2.34 ± 0.027
PGC 65436	NGC 6958	-3.8 ± 0.3	0.045	0.75 ± 0.01	1.33	1.45	16.42 ± 0.010	18.40 ± 0.042	7.16 ± 0.14	0.23 ± 0.006	2.28 ± 0.020
PGC 05326	NGC 548	-3.7 ± 0.7	0.037	0.79 ± 0.02	0.95	1.20	18.95 ± 0.018	20.57 ± 0.069	7.03 ± 0.17	0.25 ± 0.005	2.17 ± 0.018
PGC 12922	Haro 20	-3.7 ± 0.7	0.039	0.36 ± 0.04	1.02	0.91	16.29 ± 0.021	16.67 ± 0.044	2.46 ± 0.21
PGC 13360	NGC 1389	-3.3 ± 0.4	0.011	0.79 ± 0.01	1.36	1.38	16.78 ± 0.006	18.63 ± 0.025	7.21 ± 0.13	...	2.14 ± 0.026
PGC 02253	NGC 179	-3.0 ± 0.6	0.020	0.80 ± 0.01	0.97	0.99	18.38 ± 0.054	19.69 ± 0.162	6.55 ± 0.21	0.28 ± 0.008	2.38 ± 0.033
PGC 04042	NGC 392	-3.0 ± 1.2	0.058	0.79 ± 0.01	1.08	1.17	17.85 ± 0.014	18.85 ± 0.044	6.36 ± 0.14	0.29 ± 0.004	2.41 ± 0.012
PGC 05305	NGC 541	-3.0 ± 0.4	0.045	0.79 ± 0.03	1.25	1.34	17.64 ± 0.010	18.87 ± 0.032	6.97 ± 0.16	0.31 ± 0.005	2.32 ± 0.010
PGC 05311	NGC 543	-3.0 ± 0.8	0.040	0.88 ± 0.04	0.77	1.02	18.94 ± 0.018	19.80 ± 0.048	6.90 ± 0.21	0.32 ± 0.010	2.38 ± 0.014
PGC 05468	NGC 568	-3.0 ± 0.5	0.019	0.85 ± 0.03	1.35	1.35	17.96 ± 0.011	19.09 ± 0.032	6.60 ± 0.18	...	2.32 ± 0.025
PGC 05565	-3.0 ± 0.8	0.018	0.89 ± 0.03	1.16	1.08	19.12 ± 0.073	20.20 ± 0.190	6.93 ± 0.28
PGC 06957	NGC 703	-3.0 ± 1.2	0.091	0.80 ± 0.03	1.08	0.83	18.73 ± 0.104	19.21 ± 0.155	6.28 ± 0.22	0.31 ± 0.008	2.44 ± 0.029
PGC 07763	NGC 794	-3.0 ± 1.1	0.063	0.90 ± 0.03	1.11	0.96	18.43 ± 0.086	19.10 ± 0.140	6.61 ± 0.21	...	2.44 ± 0.033
PGC 10123	NGC 1023	-3.0 ± 0.3	0.061	0.82 ± 0.01	1.94	1.77	14.87 ± 0.003	16.40 ± 0.016	7.28 ± 0.06	0.27 ± 0.002	2.31 ± 0.009
PGC 10302	NGC 1060	-3.0 ± 1.0	0.209	0.86 ± 0.01	1.36	1.24	16.78 ± 0.016	17.63 ± 0.060	6.51 ± 0.15	0.31 ± 0.006	2.48 ± 0.024
PGC 11774	NGC 1222	-3.0 ± 1.2	0.060	0.42 ± 0.01	1.04	1.11	14.98 ± 0.013	15.58 ± 0.029	3.31 ± 0.13
PGC 13344	NGC 1387	-3.0 ± 0.3	0.013	0.86 ± 0.01	1.45	1.49	15.57 ± 0.003	16.54 ± 0.008	5.93 ± 0.10
PGC 13470	NGC 1400	-3.0 ± 0.3	0.065	0.78 ± 0.01	1.36	1.40	15.86 ± 0.005	16.72 ± 0.015	6.00 ± 0.13	0.30 ± 0.004	2.41 ± 0.010
PGC 21896	UGC 4014	-3.0 ± 1.3	0.025	0.66 ± 0.03	0.92	0.88	17.60 ± 0.036	18.42 ± 0.090	4.94 ± 0.22
PGC 24707	NGC 2656	-3.0 ± 0.8	0.026	1.00 ± 0.04	1.12	1.05	18.95 ± 0.017	19.97 ± 0.048	6.34 ± 0.21
PGC 29846	NGC 3163	-3.0 ± 0.7	0.013	0.89 ± 0.02	1.17	0.98	18.96 ± 0.087	19.92 ± 0.166	6.69 ± 0.22	...	2.33 ± 0.052
PGC 41538	NGC 4503	-3.0 ± 0.4	0.050	0.81 ± 0.01	1.55	1.57	16.42 ± 0.007	18.05 ± 0.016	7.18 ± 0.08	...	2.05 ± 0.090
PGC 41634	IC 3481	-3.0 ± 0.6	0.056	0.86 ± 0.02	0.92	1.22	18.45 ± 0.024	19.28 ± 0.029	6.15 ± 0.15
PGC 41939	NGC 4546	-3.0 ± 0.3	0.034	0.83 ± 0.02	1.52	1.62	15.61 ± 0.005	17.39 ± 0.018	7.22 ± 0.13	...	2.38 ± 0.048
PGC 43981	NGC 4798	-3.0 ± 1.2	0.012	0.90 ± 0.02	1.08	1.22	18.43 ± 0.014	19.82 ± 0.050	6.73 ± 0.21
PGC 51270	NGC 5574	-3.0 ± 0.9	0.031	0.69 ± 0.01	1.21	1.36	16.98 ± 0.007	19.35 ± 0.041	7.10 ± 0.14	...	1.88 ± 0.020
PGC 66113	-3.0 ± 0.5	0.032	0.64 ± 0.01	1.16	1.10	17.48 ± 0.037	19.09 ± 0.128	6.34 ± 0.18
PGC 67303	NGC 7117	-3.0 ± 0.6	0.028	0.82 ± 0.02	1.15	0.96	18.69 ± 0.063	19.74 ± 0.167	7.08 ± 0.21	0.29 ± 0.010	2.41 ± 0.055
PGC 71518	NGC 7675	-3.0 ± 0.6	0.058	0.92 ± 0.03	0.82	0.99	19.16 ± 0.022	20.75 ± 0.078	6.11 ± 0.17	...	2.26 ± 0.036
PGC 13354	NGC 1382	-2.7 ± 0.6	0.017	0.81 ± 0.02	1.18	1.55	17.38 ± 0.006	18.92 ± 0.027	7.04 ± 0.13	0.17 ± 0.035	1.98 ± 0.065
PGC 67318	NGC 7118	-2.7 ± 0.6	0.030	0.83 ± 0.01	1.17	1.06	18.36 ± 0.066	19.07 ± 0.152	6.65 ± 0.20	0.26 ± 0.010	2.28 ± 0.055
PGC 02359	NGC 193	-2.5 ± 0.6	0.023	0.90 ± 0.02	1.16	1.15	18.03 ± 0.045	19.11 ± 0.122	7.01 ± 0.19
PGC 13445	NGC 1403	-2.5 ± 0.6	0.025	0.77 ± 0.01	1.13	1.21	17.86 ± 0.057	19.23 ± 0.129	6.61 ± 0.18	0.27 ± 0.007	2.22 ± 0.014
PGC 28974	NGC 3073	-2.5 ± 0.6	0.010	0.54 ± 0.02	1.11	1.22	16.70 ± 0.006	17.50 ± 0.015	4.18 ± 0.14	...	1.54 ± 0.046
PGC 40484	NGC 4379	-2.5 ± 0.5	0.023	0.77 ± 0.01	1.28	1.29	16.83 ± 0.007	18.34 ± 0.025	6.73 ± 0.07	...	2.04 ± 0.016
PGC 60025	NGC 6359	-2.5 ± 0.6	0.021	0.78 ± 0.02	1.09	1.22	17.74 ± 0.008	18.95 ± 0.035	6.42 ± 0.14	0.28 ± 0.009	2.27 ± 0.046
PGC 68826	IC 1445	-2.5 ± 0.5	0.041	0.68 ± 0.03	1.21	1.06	18.21 ± 0.057	20.68 ± 0.271	8.19 ± 0.34
PGC 69964	NGC 7404	-2.5 ± 0.9	0.012	0.73 ± 0.02	1.17	1.25	17.56 ± 0.042	19.85 ± 0.161	8.05 ± 0.21
PGC 03367	NGC 307	-2.3 ± 0.7	0.043	0.81 ± 0.02	1.20	0.90	18.33 ± 0.057	19.94 ± 0.192	7.34 ± 0.23	...	2.49 ± 0.020
PGC 14375	NGC 1510	-2.3 ± 0.7	0.011	0.32 ± 0.01	1.12	1.61	14.63 ± 0.002	14.89 ± 0.004	1.95 ± 0.11
PGC 41781	NGC 4528	-2.3 ± 0.7	0.045	0.75 ± 0.01	1.22	1.31	17.01 ± 0.008	18.92 ± 0.024	7.04 ± 0.09	...	2.07 ± 0.026
PGC 43895	NGC 4789	-2.3 ± 0.7	0.008	0.89 ± 0.03	1.28	1.33	18.07 ± 0.012	19.22 ± 0.037	7.19 ± 0.16	0.29 ± 0.004	2.43 ± 0.008
PGC 02603	IC 48	-2.2 ± 0.7	0.038	0.69 ± 0.01	0.99	1.10	16.48 ± 0.026	16.83 ± 0.048	3.86 ± 0.14	...	2.17 ± 0.041

TABLE 1—Continued

PGC (1)	Alt. Name (2)	T (3)	$E(B - V)$ (4)	$(B - V)_{rc}$ (5)	$\log D_{25}$ (6)	$\log D_{NUV}$ (7)	NUV (8)	FUV (9)	FUV $- V_T$ (10)	Mg ₂ (11)	$\log \sigma_V$ (12)
PGC 72173	NGC 7736	-2.2 ± 0.5	0.028	0.84 ± 0.02	1.22	1.11	18.47 ± 0.081	19.82 ± 0.175	7.05 ± 0.25
PGC 00243	NGC 7808	-2.0 ± 1.2	0.043	0.69 ± 0.03	1.10	1.15	16.86 ± 0.012	17.46 ± 0.028	5.01 ± 0.14
PGC 02397	NGC 204	-2.0 ± 0.8	0.023	0.77 ± 0.02	1.08	1.04	18.22 ± 0.049	19.66 ± 0.155	6.88 ± 0.25
PGC 02959	NGC 273	-2.0 ± 0.4	0.057	0.76 ± 0.02	1.34	1.18	18.02 ± 0.069	18.84 ± 0.125	6.13 ± 0.19
PGC 03434	NGC 311	-2.0 ± 0.8	0.066	0.81 ± 0.02	1.18	1.09	18.38 ± 0.044	19.21 ± 0.098	6.45 ± 0.16	...	2.45 ± 0.024
PGC 04793	NGC 471	-2.0 ± 0.9	0.051	0.68 ± 0.01	1.02	0.93	17.21 ± 0.009	18.34 ± 0.026	5.21 ± 0.13	...	2.04 ± 0.204
PGC 04801	NGC 474	-2.0 ± 0.3	0.034	0.71 ± 0.02	1.85	1.65	16.23 ± 0.005	17.99 ± 0.021	6.63 ± 0.13	...	2.21 ± 0.013
PGC 06570	NGC 670	-2.0 ± 0.6	0.072	0.68 ± 0.01	1.31	1.19	15.44 ± 0.016	15.79 ± 0.032	3.34 ± 0.13	0.19 ± 0.010	2.12 ± 0.050
PGC 06782	NGC 687	-2.0 ± 0.8	0.062	0.86 ± 0.01	1.15	1.09	18.06 ± 0.076	19.49 ± 0.157	7.47 ± 0.20	0.30 ± 0.004	2.38 ± 0.011
PGC 08258	NGC 842	-2.0 ± 0.6	0.024	0.81 ± 0.01	1.09	1.34	17.82 ± 0.010	19.40 ± 0.037	6.86 ± 0.16
PGC 12651	NGC 1316	-2.0 ± 0.3	0.021	0.75 ± 0.01	2.08	2.09	13.58 ± 0.001	15.32 ± 0.006	6.90 ± 0.08	0.23 ± 0.015	2.36 ± 0.008
PGC 13318	NGC 1380	-2.0 ± 0.6	0.017	0.80 ± 0.01	1.68	1.73	15.29 ± 0.002	16.78 ± 0.009	6.94 ± 0.10	0.25 ± 0.008	2.35 ± 0.011
PGC 13335	NGC 1380A	-2.0 ± 0.8	0.015	0.77 ± 0.01	1.38	1.47	17.26 ± 0.006	19.21 ± 0.029	6.89 ± 0.13	...	1.83 ± 0.030
PGC 13377	NGC 1383	-2.0 ± 0.6	0.072	0.79 ± 0.01	1.28	1.21	17.80 ± 0.013	19.34 ± 0.063	7.14 ± 0.15	...	2.25 ± 0.030
PGC 13444	NGC 1394	-2.0 ± 0.8	0.073	0.82 ± 0.02	1.13	1.20	17.64 ± 0.012	18.81 ± 0.042	6.27 ± 0.14
PGC 13467	NGC 1402	-2.0 ± 0.8	0.064	0.61 ± 0.02	0.91	0.91	17.12 ± 0.009	18.20 ± 0.028	4.88 ± 0.13
PGC 14765	NGC 1553	-2.0 ± 0.3	0.013	0.75 ± 0.01	1.65	1.84	14.53 ± 0.002	16.18 ± 0.010	6.86 ± 0.08	0.26 ± 0.005	2.25 ± 0.010
PGC 25950	NGC 2784	-2.0 ± 0.3	0.214	0.81 ± 0.01	1.74	1.50	15.08 ± 0.006	16.20 ± 0.038	6.75 ± 0.14	0.29 ± 0.004	...
PGC 27765	NGC 2950	-2.0 ± 0.3	0.017	0.76 ± 0.02	1.43	1.25	16.38 ± 0.022	18.20 ± 0.079	7.35 ± 0.15
PGC 38742	NGC 4150	-2.0 ± 0.4	0.018	0.66 ± 0.01	1.37	1.30	16.17 ± 0.007	17.89 ± 0.025	6.35 ± 0.13	0.09 ± 0.005	1.93 ± 0.041
PGC 40898	NGC 4435	-2.0 ± 0.3	0.030	0.79 ± 0.01	1.44	1.79	15.55 ± 0.003	17.12 ± 0.010	6.45 ± 0.05	0.19 ± 0.003	2.19 ± 0.016
PGC 41260	NGC 4477	-2.0 ± 0.4	0.032	0.81 ± 0.01	1.58	1.63	15.74 ± 0.003	16.98 ± 0.010	6.70 ± 0.05	...	2.27 ± 0.034
PGC 41302	NGC 4479	-2.0 ± 0.6	0.029	0.81 ± 0.01	1.19	1.35	17.44 ± 0.007	18.84 ± 0.022	6.53 ± 0.06	0.18 ± 0.011	1.92 ± 0.025
PGC 42149	NGC 4578	-2.0 ± 0.5	0.021	0.72 ± 0.02	1.52	1.23	17.27 ± 0.046	18.63 ± 0.144	7.22 ± 0.16	0.28 ± 0.013	2.08 ± 0.033
PGC 53201	NGC 5770	-2.0 ± 0.8	0.038	0.73 ± 0.01	1.23	1.31	16.99 ± 0.005	19.20 ± 0.041	7.10 ± 0.14	...	2.07 ± 0.127
PGC 65376	NGC 6963	-2.0 ± 0.9	0.095	0.79 ± 0.04	0.82	1.00	18.58 ± 0.015	19.62 ± 0.064	5.96 ± 0.21
PGC 68606	NGC 7249	-2.0 ± 0.7	0.021	0.86 ± 0.04	1.05	0.81	19.91 ± 0.105	20.93 ± 0.263	7.65 ± 0.31
PGC 68612	NGC 7252	-2.0 ± 0.5	0.030	0.51 ± 0.01	1.29	1.46	15.22 ± 0.005	16.48 ± 0.017	5.21 ± 0.13	...	2.20 ± 0.053
PGC 71478	NGC 7671	-2.0 ± 0.6	0.073	0.81 ± 0.01	1.14	1.00	18.21 ± 0.081	19.52 ± 0.178	6.99 ± 0.22
PGC 71554	NGC 7679	-2.0 ± 0.4	0.065	0.36 ± 0.01	1.13	1.08	14.71 ± 0.012	15.39 ± 0.027	2.71 ± 0.13
PGC 71565	NGC 7683	-2.0 ± 0.8	0.070	0.87 ± 0.01	1.28	1.15	18.03 ± 0.080	19.45 ± 0.172	7.17 ± 0.26
PGC 13425	NGC 1393	-1.9 ± 0.4	0.067	0.76 ± 0.02	1.29	1.24	17.38 ± 0.010	18.63 ± 0.041	6.86 ± 0.14	...	2.06 ± 0.030
PGC 13091	NGC 1352	-1.8 ± 0.6	0.038	0.78 ± 0.02	1.02	1.00	18.85 ± 0.102	20.94 ± 0.280	7.77 ± 0.31
PGC 39922	NGC 4292	-1.7 ± 0.5	0.021	0.74 ± 0.02	1.22	1.37	16.99 ± 0.007	18.48 ± 0.025	6.38 ± 0.15	...	1.76 ± 0.022
PGC 13321	NGC 1381	-1.6 ± 0.5	0.013	0.81 ± 0.01	1.43	1.46	16.75 ± 0.005	18.31 ± 0.018	6.90 ± 0.10	0.25 ± 0.020	2.18 ± 0.011

NOTES.—Cols. (1) and (2): PGC number and alternate name. Col. (3): Morphological type from RC3. Col. (4): Color excess in magnitude from the Schlegel et al. (1998) reddening maps, used for the Galactic extinction corrections. Col. (5): Total (asymptotic) color index from RC3, converted into the AB system and corrected for Galactic extinction. Col. (6): Decimal logarithm of the apparent major isophotal diameter from RC3. Unit of D_{25} is $0.1''$. Col. (7): Decimal logarithm of the major diameter of the elliptical aperture determined by the *GALEX* pipeline in the NUV image. Unit of D_{NUV} is $0.1''$. Col. (8): NUV magnitude (in the AB system) given by the *GALEX* pipeline in the elliptical aperture D_{NUV} (col. [7]) and corrected for Galactic extinction. Col. (9): FUV magnitude (in the AB system) in the aperture D_{NUV} used for the NUV magnitude integration and corrected for Galactic extinction. Col. (10): FUV $- V_T$ color calculated with FUV $- V_T = \text{FUV} + (B - V)_{rc} - B_{rc}$, where B_{rc} is the B_T from RC3 converted into the AB system and corrected for Galactic extinction. Col. (11): Mg₂ index from the compilation of Golev & Prugniel (1998). Col. (12): Decimal logarithm of the central velocity dispersion σ_V in km s^{-1} from the LEDA database.

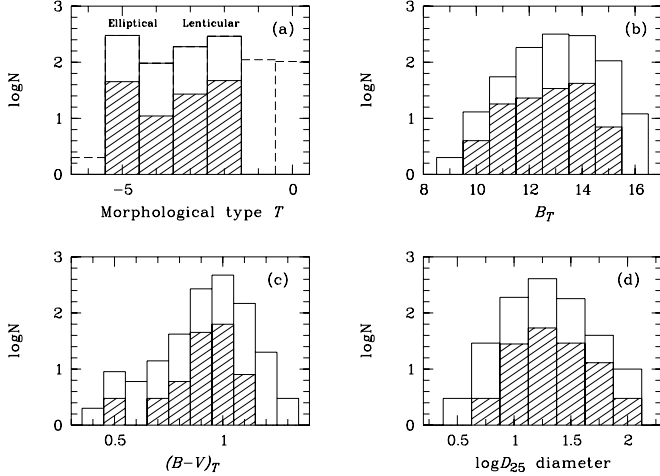


FIG. 1.— Characteristics of our sample. Shown are the numbers of galaxies (in $\log N$) in our sample (hatched histograms) compared to the early-type galaxies ($-5.5 < T < -1.5$) in the RC3 catalog entries having $(B - V)_T$ measurements in bins of (a) morphological type T , (b) B_T magnitudes, (c) $(B - V)_T$ colors, and (d) D_{25} diameter. The values of B_T and $(B - V)_T$ are in the magnitude system adopted in the RC3.

Our sample includes $\sim 15\%$ of the RC3-selected early types, without strong biases in the distribution of morphological types, B_T magnitudes, $(B - V)_T$ color, or D_{25} diameter (Fig. 1). We conclude that our sample is relatively well representative of the early-type galaxies in the RC3.

The apparent magnitudes were corrected for galactic extinction using the Schlegel et al. (1998) reddening maps and assuming the extinction law of Cardelli et al. (1989). The resulting ratio of the extinction in the *GALEX* bands to the color excess $E(B - V)$ are $A_{FUV}/E(B - V) = 8.376$ and $A_{NUV}/E(B - V) = 8.741$ (Wyder et al. 2005). The mean color excess due to the Galactic extinction is $E(B - V) \simeq 0.04$. The optical color $(B - V)_T$ is converted into the AB system and corrected for Galactic extinction by $(B - V)_{Tc} = (B - V)_T - 0.119 - E(B - V)$. Since the median redshift of our sample is $z \sim 0.01$, no K -corrections are applied.

3. ANALYSIS

3.1. FUV – NUV versus $(B - V)_{Tc}$ Color-Color Diagram

We have chosen to display our data in a FUV – NUV versus $(B - V)_{Tc}$ color-color diagram (Fig. 2). This choice, instead of a color-magnitude diagram, allows us to avoid significant dispersion due to uncertainties in distance (our sample of nearby galaxies spans a redshift range $0 < z < 0.045$). The choice of the FUV – NUV color, based on *GALEX* measurements alone, minimizes instrumental inhomogeneities. The FUV – NUV color is calculated using the FUV and NUV magnitudes (Table 1) measured by the *GALEX* pipeline through the elliptical aperture whose major axis is scaled to 2.5 times the Kron diameter determined in the NUV images. We emphasize that the UV color is calculated in the same aperture. The major diameter D_{NUV} of this aperture (Table 1) is similar to the D_{25} values given in the RC3.

Our sample of early-type galaxies is divided into 56 ellipticals (defined as galaxies with morphological type $-5.5 \leq T < -3.5$) and 74 lenticulars (defined as galaxies with morphological type $-3.5 \leq T < -1.5$). The color-color diagram clearly shows a tight anticorrelation between the FUV – NUV and $(B - V)_{Tc}$ colors for elliptical galaxies: the UV color decreases when $(B - V)_{Tc}$ increases (Fig. 2a). The amplitude range in the UV color is 2 mag, against only 0.2 mag in $(B - V)_{Tc}$. When

the two galaxies off the correlation are excluded, the correlation coefficient is $r = -0.73$. This value reaches -0.84 if we select ellipticals with a mean error on $(B - V)_{Tc} < 0.02$. For the early-type galaxies with morphological types in the range $-3.5 < T < -1.5$ that are lenticular galaxies (Fig. 2b), the relation is more dispersed, and a larger fraction of objects are off the correlation.

Because elliptical galaxies exhibit a FUV optical color gradient where the inner regions are slightly bluer than the outer parts (e.g., Ohl et al. 1998; Rhee et al. 2005), we have examined how this gradient may affect the UV color in different apertures. The *GALEX* pipeline allows us to calculate the mean FUV – NUV color in seven circular apertures with diameters fixed at $3.0''$, $4.5''$, $7.5''$, $12.0''$, $18.0''$, $25.5''$, and $34.5''$ and through two elliptical apertures whose major-axis D_{NUV} and D_{FUV} are scaled to 2.5 times the Kron diameter determined in the NUV and FUV images, respectively. As expected, the mean FUV – NUV color becomes bluer when the aperture radius decreases. In our sample, the difference between the mean color in the $34.5''$ circular aperture and the mean color in the D_{NUV} elliptical aperture is only 0.03 mag. There is no significant difference between the mean colors when the UV fluxes are measured through the D_{NUV} or D_{FUV} elliptical apertures.

Gil de Paz et al. (2007) have presented the integrated photometry, surface brightness, and color profile for a total of 1034 nearby galaxies observed by *GALEX*. They obtain a main difference of -0.19 ± 0.20 and -0.23 ± 0.20 between the asymptotic magnitudes and the D_{25} aperture magnitudes for the FUV and NUV, respectively, resulting in a difference of 0.04 in the FUV – NUV color. Their sample includes 48 ellipticals ($-5.5 \leq T < -3.5$, mean error on $T < 1.5$) and 49 lenticulars ($-3.5 \leq T < -1.5$, mean error on $T < 1.5$) with FUV and NUV asymptotic magnitudes and B , V measurements. Among them, 73 are in common with our sample. The comparison of our FUV and NUV with asymptotic magnitudes measured by Gil de Paz et al. shows a mean difference of -0.06 ± 0.27 and -0.19 ± 0.20 in the FUV and NUV, respectively. Our UV magnitudes are very close to the integrated aperture magnitudes in D_{25} ellipses. The mean differences are 0.03 ± 0.18 for $NUV_{D_{25}} - NUV$ and 0.06 ± 0.21 for $FUV_{D_{25}} - FUV$. The similarity of the D_{25} and D_{NUV} diameters explains the small differences between the $NUV_{D_{25}}$, $FUV_{D_{25}}$ aperture magnitudes and those used in this paper.

We conclude that the FUV – NUV colors obtained with large apertures or asymptotic photometry are very similar for our sample. The trends in the FUV – NUV diagrams versus $(B - V)_{Tc}$ are not significantly affected, even if the strength of the relations is slightly changed (the correlation coefficient is -0.61 for the elliptical galaxies in common with the Gil de Paz et al. [2007] sample and asymptotic magnitudes).

3.2. UV Color versus Mg_2 and Velocity Dispersion

Using *IUE* observations of early-type galaxies, Faber (1983) and Burstein et al. (1988) found that the nuclear $(1550 - V)$ colors become bluer as the nuclear Mg_2 spectral line index increases. Burstein et al. (1988) also found significant correlations between $1550 - V$ and the central velocity dispersion. More recently, using a large sample of SDSS early-type galaxies observed by *GALEX*, Rich et al. (2005) reported the lack of a significant correlation between the $(FUV - r)$ color and the Lick Mg_2 index or the velocity dispersion, while Boselli et al. (2005) found a mild trend between FUV – NUV and Mg_2 in a sample of early-type galaxies in the Virgo Cluster.

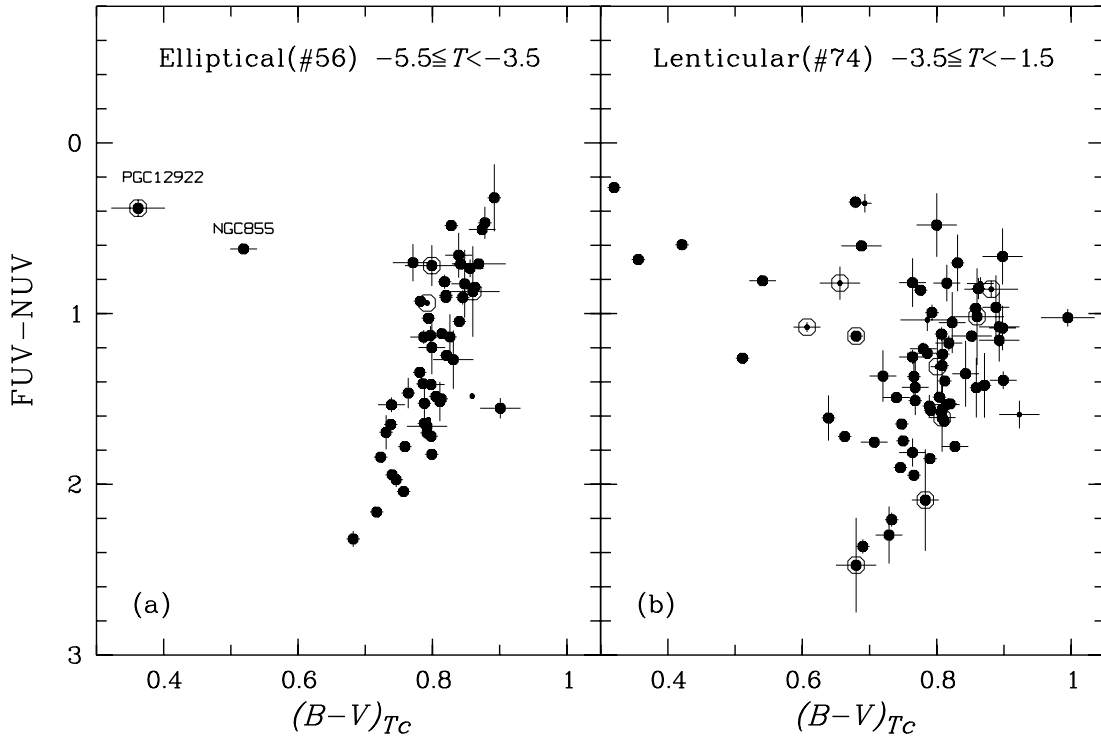


FIG. 2.—Color-color diagrams $FUV - NUV$ vs. $(B - V)_{Tc}$ for the elliptical and lenticular subsamples. The elliptical subsample (a) shows a clear anticorrelation. Two galaxies, PGC 12922 and PGC 8557 (NGC 855), are off the correlation. NGC 855 appears similar to an E6 or E7 galaxy on the POSS plate, whereas *HST* observation of the inner structure clearly shows features of an edge-on late-type galaxy (Phillips et al. 1996), and Walsh et al. (1990) find an extended H I distribution using the VLA. PGC 12922 is the peculiar blue galaxy Haro 20. The lenticular subsample (b) shows a more dispersed anticorrelation and a larger fraction of galaxies off the correlation. Open circles indicate the galaxies for which the NUV or FUV flux contribution in a central $6''$ aperture is larger than 30%. Small filled circles indicate galaxies with major isophotal diameter $\log D_{25} < 1$ (i.e., major isophotal diameter $< 1'$). We conclude that we are seeing an extension of the color-metallicity relationship for early-type galaxies into the UV (see text). The larger scatter and greater number of outliers suggests that lenticulars have a more complex star formation history.

We have plotted in Figure 3 the $FUV - NUV$ and $FUV - V_T$ colors versus the Mg_2 index and the central velocity dispersion σ_V for our sample. The Mg_2 index comes from the compilation of Golev & Prugniel (1998), and the σ_V comes from the LEDA database.¹⁵

Figure 3 (top) shows clearly that the $FUV - NUV$ color of ellipticals becomes bluer as the spectral line index Mg_2 increases. The same trend appears between the UV color and the central velocity dispersion σ_V . It should be noted that the Mg_2 and σ_V values given by Golev & Prugniel (1998) and the LEDA database are standardized to an aperture of $0.595 h^{-1}$ kpc (equivalent to an angular diameter of $3.4''$ at the distance of the Coma Cluster). We have checked that when the $FUV - NUV$ colors are calculated in a similar aperture (by linear interpolation between the UV flux measured by the *GALEX* pipeline through seven circular apertures), the trends in Figure 3 (top) are not significantly affected. These correlations are the opposite of the well-known dependence of $B - V$ color on metal abundance or central velocity dispersion and are not surprising, since, as shown in § 3.1, there is an anticorrelation between the $FUV - NUV$ and $B - V$ colors. If we consider that the velocity is a mass indicator, the more massive ellipticals tend to be bluer in $FUV - NUV$. If the velocity dispersion also correlates with the galaxy age (Thomas et al. 2005; Ziegler et al. 2005), the bluest in UV color would be the oldest. Lenticulars show a similar trend, but a larger fraction of galaxies are off the correlation.

The correlations found above with the $FUV - NUV$ color (Fig. 3, top) for the subsample of morphological ellipticals be-

come marginal (not visible against velocity dispersion) with the $FUV - V_T$ color and suffer far greater scatter (Fig. 3, bottom). For instance, the correlation coefficient with the Mg_2 index decreases from -0.83 to -0.63 (NGC 855 and NGC 6482 excluded). This is expected because the $FUV - V_T$ color is not as homogeneous in terms of instrument and photometric aperture as the $FUV - NUV$ color. The correlations with the $FUV - V_T$ color are not visible for the lenticular subsample, suggesting that the selection of the sample is, in addition to the homogeneity of the UV color, a factor in the strength of the correlation.

For the morphological ellipticals only, we recover an approximation of the correlation between $FUV - V_T$ and Mg_2 index that was seen by Burstein et al. (1988) in *IUE* data. However, Burstein et al. also saw a correlation with velocity dispersion that we do not recover here, and the lenticular subset also fails to show a correlation. In comparison with Rich et al. (2005), our elliptical subset has been selected to be uniform in morphology, exhibits a tight color-color correlation, and covers a wide range in the Mg_2 index. These factors may have favored our detection of the mild correlation we report here.

3.3. Discussion

Before any interpretation of the color-color diagrams of Figure 2, we have evaluated the possible UV flux contamination from AGN activity by measuring the contribution of the central part of the galaxies to the UV emission. This evaluation is possible because the major part of the galaxies in our sample are spatially resolved relatively well by *GALEX*. We have estimated the UV flux in the central $6''$ aperture (the typical angular resolution of *GALEX*) by linear interpolation between the UV flux measured by the *GALEX* pipeline through the $4.5''$ and $7.5''$

¹⁵ Available at <http://leda.univ-lyon1.fr>.

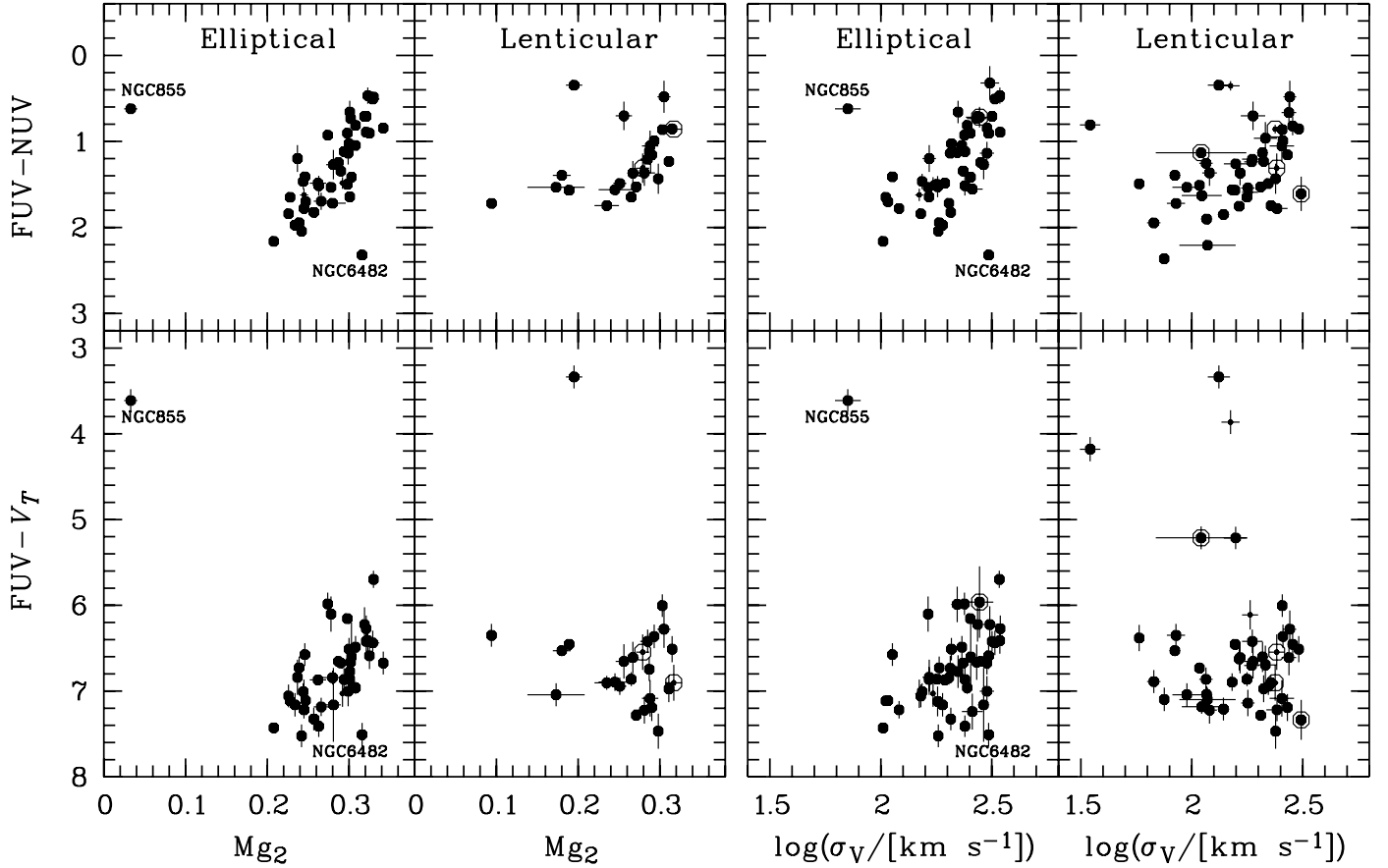


FIG. 3.—FUV – NUV and FUV – V_T colors vs. Mg_2 index and central velocity dispersion σ_V . The Mg_2 and σ_V values come from the compilation of Golev & Prugniel (1998) and the LEDA database, respectively. The elliptical subsample shows an anticorrelation (FUV – NUV color decreases when Mg_2 or σ_V increases) which is reversed in sense from the well-known dependence of $B - V$ color on metal abundance or central velocity dispersion. This correlation gets worse with either the FUV – V_T color or the lenticular subsample. The symbols are the same as in Fig. 2.

apertures. Among the 119 galaxies with $\log D_{25} > 1$ (i.e., major apparent diameter $> 60''$), we find 3 galaxies in the NUV band and 8 in the FUV band for which the central ($< 6''$) UV flux (stellar population + AGN) is $> 30\%$ of the total UV flux. These galaxies, likely candidates for AGNs, are indicated in Figure 2 (*open circles*). The trends in the color-color diagrams are not affected by the possible AGN activity.

From Figure 2 it appears that morphological elliptical galaxies follow a much tighter anticorrelation between FUV – NUV and $B - V$ than the lenticular galaxies. We propose that we are seeing a metallicity-driven correlation, with the greater dispersion for the lenticulars arising from residual star formation. We suspect that blanketing in the NUV band (consider this band an extended U band) is the primary driver of this observed correlation. The light in the NUV bandpass is dominated by the hottest stars in an old stellar population, which would be turnoff and subgiant stars, or blue HB stars if present. Any young stellar population would make the FUV – NUV and $B - V$ colors simultaneously bluer, according to their current spectral energy distribution. It is also conceivable that factors such as age, metallicity, mass, or a combination of these, currently advocated for the spread in the $B - V$ colors of early-type galaxies, may redden the $B - V$ color while simultaneously decreasing the NUV and/or increasing the FUV radiation and making the FUV – NUV bluer. Figure 4 clearly shows a positive correlation between the NUV – V_T and $B - V$ colors and a marginal anticorrelation between the FUV – V_T and $B - V$ colors for ellipticals. This is consistent with our suggestion that the stars contributing to the color-color correlation are

likely old metal-rich stars, with the dominant contributors in the NUV band being the turnoff and subgiant stars (e.g., Dorman et al. 2003). The tight correlation for the ellipticals appears consistent with little intrinsic variation, from galaxy to galaxy, of the hot stellar populations (e.g., post-AGB or EHB stars) dominating the FUV band. The dominant driver in this color-color relationship appears to be metallicity and (in the case of lenticulars) star formation.

It is of interest to check how these results compare to the predictions of the stochastic library of model star formation histories used by Salim et al. (2005) to interpret the seven-band photometry of the combined sample of *GALEX* and SDSS galaxies. We note that this library is based on the Bruzual & Charlot (2003) population synthesis models, which were not optimized to describe the UV upturn of early-type galaxies (in these models, the UV light of old stellar populations is produced primarily by post-AGB rather than EHB stars). First, the anticorrelation between the FUV – NUV and $B - V$ colors is very well reproduced (Fig. 5a) if we select models representative of old galaxies (mean age of star formation episodes in the 8–12 Gyr interval) with low extinction ($\tau_V < 0.5$). These galaxies formed the bulk of their stars at $z > 1$ (assuming $\Omega_M = 0.3$, $\Omega_\Lambda = 0.7$, $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$). The spread in UV and $B - V$ colors can be easily reproduced by the variation of metallicity, galaxies with high metallicity values ($1.0 < Z/Z_\odot < 2$) being bluer in UV color (FUV – NUV $\lesssim 1.2$), while the galaxies with low metallicity ($0.2 < Z/Z_\odot < 1.0$) are redder (FUV – NUV $\gtrsim 1.2$). The models have old stellar populations, but the only hot star

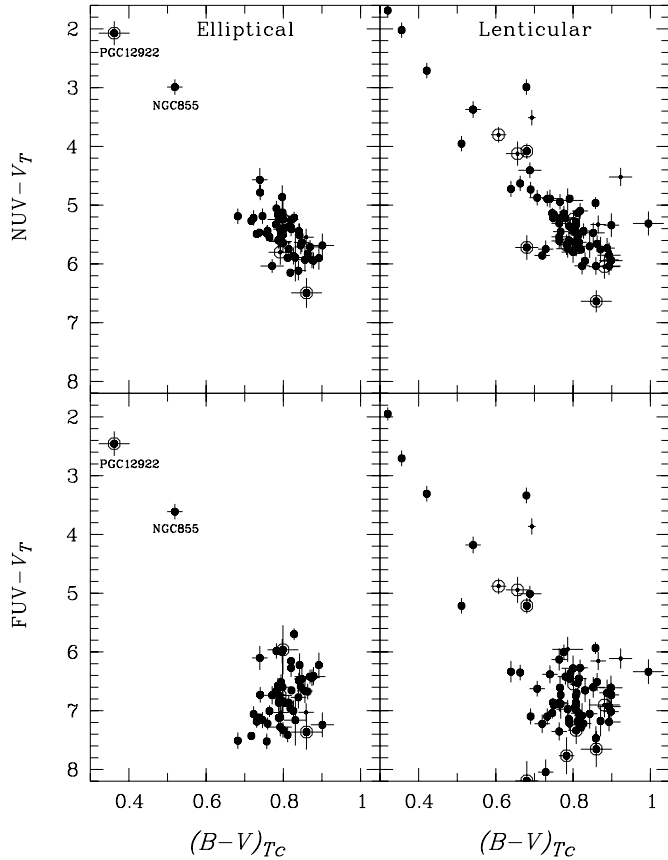


FIG. 4.—Color-color diagrams $NUV - V_T$ and $FUV - V_T$ vs. $(B - V)_{Tc}$. The symbols are the same as in Fig. 2.

contributors are post-AGB stars which do not have a metallicity dependence in the models and are not the contributors favored by the most recent observations (Brown et al. 2000; Brown 2004). This suggests again that blanketing in NUV is presumably the main cause of the agreement between the models and the UV color-color correlation for ellipticals. The FUV flux would therefore be dominated by hot stars (whose post-AGB in the models is only proxy) that would have low blanketing, as suggested by the spectra of Brown et al. (1997). The stars providing most of the NUV flux would be those near the main-sequence turnoff.

Second, the departure from the trend (or the simultaneous decrease of the $FUV - NUV$ and $B - V$ colors) is well reproduced with the models with a mean age of star formation episodes in the 4–12 Gyr interval (Fig. 5b) and can be interpreted as recent star formation activity. We expect low impact of the internal extinction in the scatter in Figure 5 because the internal extinction is generally very small in most nearby early-type (Dorman et al. 2003) galaxies, and the $FUV - NUV$ color is essentially reddening-free for moderate amounts of reddening. In fact, the color-color correlation may prove to be a powerful detector of weak star formation activity. Figure 2 shows that it may be the case for 4% of the ellipticals and about 15% of the lenticulars. This would mean about 10% of early-type galaxies have residual star formation in our full sample of 130 early-type galaxies. This number is lower than the fraction of 15% reported by Yi et al. (2005) with a different approach based on a NUV color-magnitude relation and the sample of galaxies identified as early type by Bernardi et al. (2003) in the SDSS. That the derivation of this number depends on the selection criteria used to define the sample of early-type galaxies is shown directly by our approach; the fraction of early-type galaxies with signs of star

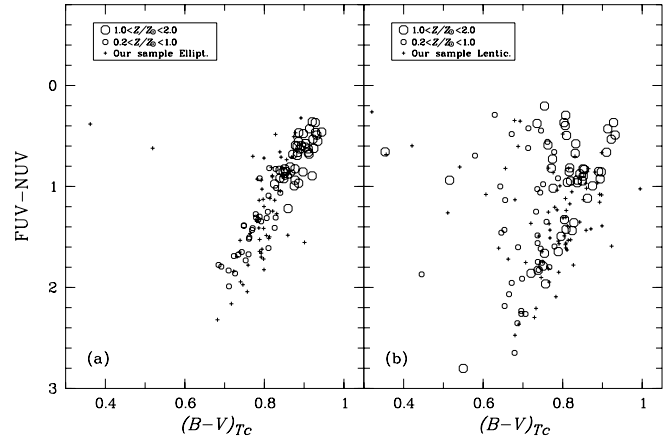


FIG. 5.—Comparison of the observed colors (crosses) as displayed in Fig. 2 with the stochastic realizations of model star formation histories as generated by Salim et al. (2005) from the Bruzual & Charlot (2003) population synthesis models (open circles). (a) Stochastic realizations representative of old galaxies: exponentially declining star formation law $[SFR(t) \propto \exp(-\gamma t)]$ with $\gamma > 0.9$, mean age of star formation episodes in the 8–12 Gyr interval, and low extinction (effective V -band absorption optical depth $\tau_V < 0.5$). (b) Same selection but with a mean age of star formation episodes in the 4–12 Gyr interval. The models well reproduce both the tight anticorrelation between the UV and $B - V$ colors for the elliptical subsample (a) and the scatter for the lenticular subsample (b).

formation increases from 4% to 15% when selecting the RC3 morphological types $-3.5 < T < -1.5$ instead of $-5.5 < T < -3.5$. It is likely that a number of the features leading to a non-classification in the types $-5.5 < T < -3.5$ are somewhat related to star formation activity. At large distances, selection of early-type objects based on concentration index, luminosity profile, and spectra may naturally lead to a larger fraction of residual star formation. Conversely, it should be possible to combine a color selection with the morphological selection of nearby galaxies that would retain only passively evolving early-type galaxies and no residual star formation.

Another illustration of how the fraction of early-type galaxies with residual star formation depends on the selection of the objects is shown by Boselli et al. (2005); their sample of early-type galaxies in the Virgo Cluster contains a large fraction of dwarf galaxies whose colors, different from those of giant ellipticals, are interpreted in terms of recent star formation activity.

4. CONCLUSIONS

From the cross-match of the *GALEX* observations with the RC3 catalog, we have obtained a reliable sample of 130 early-type galaxies, divided in two subsamples: 56 ellipticals ($-5.5 \leq T < -3.5$) and 74 lenticulars ($-3.5 \leq T < -1.5$). The main results are as follows:

1. There is an anticorrelation between the $FUV - NUV$ and $(B - V)_{Tc}$ colors (the UV color decreases by 2 mag when $(B - V)_{Tc}$ increases by 0.2 mag) in the elliptical subsample. This relation is more dispersed, and a larger fraction of objects are off the correlation in the lenticular subsample. We conclude that this correlation reflects an extension into the UV of the color-metallicity relationship for ellipticals and is mainly driven by metal line blanketing in the NUV band. Most of the flux in this band would originate from stars near the main-sequence turnoff while the FUV flux is dominated by hot evolved stars whose precise nature has yet to be determined.
2. There is a clear correlation between the $FUV - NUV$ color and the Mg_2 spectral index (also velocity dispersion) for the elliptical subsample. However, this correlation gets worse

with either the FUV – V_T color or the lenticular subsample (to the point of disappearing with a combination of both). We conclude that this correlation is an additional aspect of the color-metallicity relationship we noted earlier.

3. The anticorrelation between FUV – NUV color and $B - V$ is especially tight for morphological ellipticals while showing more scatter for lenticulars. We find this to be a sensitive means of detecting residual star formation; 10% of our sample shows evidence for residual star formation. We are able to model this effect using the Bruzual & Charlot (2003) models.

4. The trends found either diminish in quality (less tight relations, large number of outliers) or disappear from the elliptical to the lenticular subsample. Relevant properties of early-type galaxies are therefore crucially dependent on how strictly the sample of these objects is defined.

The sample of 130 galaxies presented in this paper represents $\sim 15\%$ of the reliable early types of the RC3 catalog; when the *GALEX* survey is completed we can expect to multiply by a factor of ~ 6 the number of early-type galaxies in the sample.

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Facilities: GALEX

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